

GALACTIC ANGULAR MOMENTUM IN COSMOLOGICAL ZOOM-IN SIMULATIONS. I. DISK AND BULGE COMPONENTS AND THE GALAXY–HALO CONNECTION

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ABSTRACT

We investigate the angular momentum evolution of four disk galaxies residing in Milky Way-sized halos formed in cosmological zoom-in simulations with various sub-grid physics and merging histories. We decompose these galaxies kinematically and photometrically, into their disk and bulge components. The simulated galaxies and their components lie on the observed sequences in the j_* – M_* diagram relating the specific angular momentum and mass of the stellar component. We find that galaxies in low-density environments follow the relation $j_* \propto M_*^\alpha$ past major mergers, with $\alpha \sim 0.6$ in the case of strong feedback, when bulge-to-disk ratios are relatively constant, and $\alpha \sim 1.4$ in the other cases, when secular processes operate on shorter timescales. We compute the retention factors (i.e. the ratio of the specific angular momenta of stars and dark matter) for both disks and bulges and show that they vary relatively slowly after averaging over numerous but brief fluctuations. For disks, the retention factors are usually close to unity, while for bulges, they are a few times smaller. Our simulations therefore indicate that galaxies and their halos grow in a quasi-homologous way.

1. INTRODUCTION

The mass M and angular momentum J are two of the most basic properties of galaxies. For many purposes, it is more convenient and physically meaningful to describe galaxies in terms of their mass M and specific angular momentum $j \equiv J/M$ (for the stellar parts of galaxies, we denote these quantities by M_* and j_*). Galaxies of the same disk-to-bulge ratio or morphological type at redshift $z = 0$ obey scaling relations of the form $j_* \propto M_*^\alpha$, with $\alpha \sim 2/3$ (Fall 1983; Romanowsky & Fall 2012; Fall & Romanowsky 2013). In a plot of $\log j_*$ against $\log M_*$, disk-dominated (Sc, Sd) and bulge-dominated (E) galaxies lie along roughly parallel sequences of slope $\sim 2/3$ separated by a factor of ~ 5 in j_* at each M_* . Galaxies of intermediate types (Sb, Sa, and S0) populate the region between these sequences.

The observed sequence of disk-dominated galaxies at $z = 0$ in the j_* – M_* diagram is close to the predictions of a simple analytical model in which galactic disks have the same specific angular momenta as their dark-matter halos (Fall & Efstathiou 1980; Fall 1983; Mo et al. 1998).

Disk-dominated galaxies at high redshift ($0.2 < z < 3$) also appear to obey this simple model (Burkert et al. 2016; Contini et al. 2016). Moreover, the sizes of galactic disks, another reflection of their angular momenta, are consistent with this model, both at $z = 0$ (Kravtsov 2013) and at $0 < z < 3$ (Huang et al. 2016).

Over the past two decades, there have been many attempts to reproduce the observed j_* – M_* sequences in hydrodynamical simulations of galaxy formation. Until recently, most of these simulations produced galaxies that lay closer to the bulge-dominated sequence than to the disk-dominated sequence (e.g. Navarro & Steinmetz 1997; Weil et al. 1998; Abadi et al. 2003; Stinson et al. 2010a). The failure of the early simulations to reproduce the observed disk-dominated sequence has been called the “angular momentum problem.” It is another manifestation or close relative of the “over-cooling problem.”

The situation has changed dramatically in the past few years as a result of greater computing power, better numerical techniques, and the inclusion of more realistic physical processes in the simulations. Feedback – the injection of momentum and/or energy into the interstellar and/or circumgalactic media (ISM and CGM)

– appears to be crucial (Okamoto et al. 2005; Governato et al. 2007; Scannapieco et al. 2008; Zavala et al. 2008). At the same time, increased mass and spatial resolution have reduced numerical artefacts responsible for spurious angular momentum losses in galactic disks (Okamoto et al. 2003; Kaufmann et al. 2007; Stinson et al. 2010b). Some of the simulations are now capable of producing respectable galactic disks (e.g. Okamoto et al. 2005; Governato et al. 2007; Scannapieco et al. 2008; Guedes et al. 2011; Agertz et al. 2011; Aumer et al. 2013; Marinacci et al. 2014; Roškar et al. 2014; Murrante et al. 2015; Colín et al. 2016; Lagos et al. 2016). The most recent generation of large-volume hydrodynamical simulations of galaxy formation (e.g. Illustris and Eagle; Vogelsberger et al. 2014; Schaye et al. 2015) have also succeeded, at least approximately, in reproducing the observed sequences of disk-dominated, bulge-dominated, and intermediate-type galaxies (Genel et al. 2015; Pedrosa & Tissera 2015; Teklu et al. 2015; Zavala et al. 2016).

The evolution of the angular momentum of galaxies and their disk and bulge components, relative to their formation history, has been studied in greater detail by Scannapieco et al. (2009) and Sales et al. (2012). Some physical processes cause losses in specific angular momentum (e.g. galaxy mergers; Jesseit et al. 2009; Capelo & Dotti 2016), while others cause gains (e.g. galactic fountains; Übler et al. 2014). However, the observed j_* – M_* diagrams for disk-dominated galaxies clearly show that those gains and losses mostly cancel out, leading to an apparent, if not strict, conservation of specific angular momentum (Romanowsky & Fall 2012; Fall & Romanowsky 2013).

The large-volume hydrodynamical simulations, with typical dimensions ~ 100 Mpc, produce many thousands of galaxies, more than enough to define the j_* – M_* relations over wide ranges of mass and disk-to-bulge ratio. The price paid for these large volumes and galaxy populations, however, is relatively low mass and spatial resolution, typically $\sim 10^7 M_\odot$ and $\sim \text{kpc}$. Many of the most important physical processes, particularly those involving transport of radiation, mass, and momentum or energy, are affected strongly by the structure of the ISM and CGM on much smaller scales. To take only one example among many, the formation of star clusters and their resulting feedback occurs in some of the densest parts of the ISM, the so-called clumps, with typical masses $\sim 10^2$ – $10^6 M_\odot$ and dimensions $\sim \text{pc}$.

These complicated, and only partially understood, small-scale processes are dealt with in the hydrodynamical simulations of galaxy formation by approximate sub-grid prescriptions rather than by direct solution of the relevant dynamical equations. Given the large mismatch of scales, the use of sub-grid modules is likely

to be necessary in this field for the foreseeable future. Thus, it is important to analyze simulations with different resolution, numerical techniques, and sub-grid prescriptions, especially for star formation and feedback by both young stars and active galactic nuclei (AGN), to determine which results depend sensitively on these features and which are robust. In this respect, high-resolution zoom-in simulations of the formation and evolution of individual galaxies are a valuable complement to the large-volume simulations of galaxy populations. In particular, they allow for a more detailed and reliable study of the processes causing gains and losses of specific angular momentum. Recently, zoom-in simulations have been able to capture even subtle internal dynamical processes occurring in disks, from non-axisymmetric instabilities such as bars to radial migration of stars, and their interplay with other galaxy properties such as the age of stellar populations (Brook et al. 2011; Bird et al. 2013; Guedes et al. 2011; Guedes et al. 2013; Gabor & Bounaud 2013; Stinson et al. 2013; Bonoli et al. 2016; Spinoso et al. 2016). Thus, the latest generation of zoom-in simulations appears to capture many of the salient features of galaxy formation and evolution.

In this and a companion paper (hereafter papers I and II), we report on a study of the angular momentum in four high-resolution zoom-in simulations of galaxy formation and evolution. The focus of paper I is on the evolution of the stellar components – the disks and bulges – of these galaxies in the j_* – M_* diagram from the beginning of the simulations at redshift $z \sim 100$ all the way to the end at $z = 0$. The focus of paper II is on the evolution of the inflowing, outflowing, and circulating gas, in different ranges of density and temperature, how it gains and loses specific angular momentum, and how this accounts for the evolution of the stellar components. Running the simulations all the way to $z = 0$ is crucial, because their behavior changes in important ways at $z \sim 1$. Our study is similar in spirit to the recent analyses of high-resolution zoom-in simulations by Fiacconi et al. (2014), Danovich et al. (2015), and Agertz & Kravtsov (2016).

The plan for the remainder of this paper is the following. In Section 2, we describe our simulations. In Section 3, we decompose the galaxies into disks and bulges. In Section 4, we plot galaxies and their components in the j – M diagram at $z = 0$ and at higher redshifts. In Section 5, we study the relation between the specific angular momentum of galaxies and their dark halos. Finally, in Section 6, we discuss and summarize our conclusions. The adopted cosmological parameters in all four simulations are $\Omega_M = 0.24$, $\Omega_\Lambda = 1 - \Omega_M$, $\Omega_b = 0.042$, $H_0 = 73 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $n_s = 0.96$, and $\sigma_8 = 0.76$, based on the first three years of data from the Wilkinson Microwave Anisotropy Probe (Spergel et al.

2007).

2. SIMULATIONS

We analyze four high-resolution cosmological zoom-in simulations of Milky Way-sized galaxies. The simulations were performed with the tree-smoothed particle hydrodynamics (SPH) code GASOLINE (Wadsley et al. 2004) with mass resolution $m_{\text{dm}} \simeq 9.8 \times 10^4 M_{\odot}$ and $m_{\text{SPH}} \simeq 2 \times 10^4 M_{\odot}$, and spatial resolution $\simeq 120$ pc. The zoom-in technique (Katz & White 1993) is well established numerically after more than two decades from its introduction, but care has to be taken in building the initial conditions to avoid numerical artifacts which could affect the dynamics. For the simulations presented here, the original periodic low-resolution box from which the initial conditions were subsequently refined is much larger than the Lagrangian subvolume that was selected for the refinement. The large-scale box has indeed a size of 90 Mpc as opposed to about 1 Mpc for the Lagrangian high-resolution subvolume at $z = 0$ (different for Venus, being 60 Mpc and 0.2 Mpc, respectively). The total number of particles in the box of Eris is 53 million (including 13 million of gas), whereas the box of Venus has 170 million particles (16 million of gas).

While the base box is larger than in other published zoom-in simulations, as explained in Katz & White (1993), Mayer et al. (2008), and Governato et al. (2004), choosing a large enough box for the coarsely resolved region is important because lack of large-scale power may bias the angular momentum of collapsing halos. In building the initial conditions we also checked that the spin parameter of the selected halo remains essentially unchanged as we introduce successive refinements.

Table 1. Input parameters of the runs

Run	UVB	IMF	n_{SF}	ϵ_{SN}	MC	AGN	IC
Eris	HM96	K93	5	0.8	low-T	no	Q
Venus	HM96	K93	5	0.8	low-T	no	A
EBH	HM96	K93	5	0.8	low-T	yes	Q
E2k	HM12	K01	100	1.0	all-T	no	Q

NOTE—UVB – UV background (HM96: Haardt & Madau 1996, HM12: Haardt & Madau 2012), IMF – initial mass function (K93: Kroupa et al. 1993, K01: Kroupa 2001), n_{SF} – star formation density threshold, ϵ_{SN} – SN efficiency parameter, MC – metal cooling, AGN – AGN feedback, and IC – initial conditions (Q: quiet merger history, A: active merger history).

One of the runs, Eris, has been shown to be extremely successful in recovering various properties of late-type spirals such as the Milky Way (Guedes et al. 2011). The other runs comprise two which stem from the same initial conditions but have different sub-grid models, and

a fourth one with different initial conditions. They are, in order: Eris2k (hereafter E2k, described in more detail in Sokolowska et al. 2016) for which sub-grid parameters were tuned to yield a stronger effect of supernova (SN) feedback to lower star formation rates at high redshift; ErisBH (hereafter EBH) being a replica of Eris that includes AGN feedback and yields final correlations between galaxy properties and the mass of the central supermassive black hole that are in good agreement with those of late-type spirals (Bonoli et al. 2016); and Venus, with the same sub-grid physics as the original Eris but different initial conditions, chosen to have an active merging history down to low redshift in contrast with the quiet merging history of the other three runs but also a nearly identical final virial halo mass ($\sim 8 \times 10^{11} M_{\odot}$) and halo spin parameter ($\lambda \sim 0.03$). Some important simulation parameters of all runs, including the choice of the UV background and aspects of the sub-grid physics, are listed in Table 1 and discussed below.

All runs include radiative and Compton cooling. However, in Eris, EBH, and Venus, gas cooling is computed for a simple mixture of H and He via non-equilibrium cooling rates in the presence of the ionizing cosmic ultraviolet (UV) background (Haardt & Madau 1996; Wadsley et al. 2004). Additionally, gas of $T < 10^4$ K cools through fine structure and metastable lines of C, N, O, Fe, S, and Si (Bromm et al. 2001; Mashchenko et al. 2007). In E2k, we instead account for metal-line cooling at all temperatures, employing tabulated rates computed with the code CLOUDY (Ferland et al. 1998), which assumes that metals are in ionization equilibrium (Shen et al. 2010) in the presence of an updated cosmic ionizing background (Haardt & Madau 2012).

The recipes for star formation and SN feedback are the same in all the runs and are described in Stinson et al. (2006). Gas particles must be dense – namely have a density above the threshold n_{SF} (set to 100 atom cm^{-3} in E2k and 5 atom cm^{-3} in the other runs) – and cool ($T < T_{\text{max}} = 1\text{--}3 \times 10^4$ K) in order to form stars. Particles which fulfill these requirements are stochastically selected to form stars according to $dM_*/dt = c^* M_{\text{gas}}/t_{\text{dyn}}$, where M_* is the mass of stars created, c^* is a constant star formation efficiency factor (set to 0.1 in all runs), M_{gas} is the mass of gas creating the star, and t_{dyn} is the gas dynamical time. Each star particle then represents a population of stars, covering the entire initial mass function (IMF; listed in Table 1).

Stars more massive than $8 M_{\odot}$ explode as SNII. According to the “blastwave feedback” model of Stinson et al. (2006), feedback is purely thermal, as the blastwave shocks convert the kinetic energy of ejecta into thermal energy on scales smaller than those resolved by our simulations. Once energy is ejected (the fraction

of SN energy that couples to the interstellar medium is $\epsilon_{\text{SN}} = 1.0$ in E2k and 0.8 in the remainder), particles receiving the energy are prevented from cooling for typically 10–50 Myr, with the cooling shut-off timescale being computed as the sum of the Sedov–Taylor (Taylor 1950; Sedov 1959) and snow-plough phases in the ejecta (McKee & Ostriker 1977). By delaying the cooling, we model in a phenomenological way the unresolved effect of momentum and energy input by turbulent dissipation in the ejecta before they reach the radiative phase. The strength of feedback depends on the number of SNe produced, which in turn is governed by the IMF and, locally, by the star formation density threshold.

The IMF in Eris, EBH, and Venus is from Kroupa et al. (1993), whereas the IMF in E2k is the updated one from Kroupa (2001), which yields about a factor of 2.8 more SNe for the same star formation rate. Furthermore, as explained in detail in Guedes et al. (2011) and Mayer (2012), the local star formation rate, and thus the local effect of SNe, can be boosted significantly by raising the star formation density threshold as the interstellar medium is allowed to become more inhomogeneous, an effect that saturates only at very high resolution and density thresholds, well above those resolved with cosmological simulations (Hopkins et al. 2012). This implies that in E2k heating by SN feedback is boosted both globally and locally. We recall that E2k is a run that follows an extensive study of sub-grid parameters by running many different simulations with the same Eris-type initial conditions in order to determine the combination of parameters that yields realistic stellar masses in accordance with abundance matching at both high and low redshift, these being shown in Table 1. Indeed in the original Eris suite the conversion of gas into stars was too efficient at high redshift, although final stellar masses at $z = 0$ are in agreement with abundance matching (see also Agertz & Kravtsov 2015; Sokolowska et al. 2016). E2k also has a richer inventory of physical processes, not only metal-line cooling but also a sub-grid turbulent diffusion prescription for both metals and thermal energy which allows mixing to be captured in SPH (Shen et al. 2010). In all the runs, metals come from SNI and SNII (Stinson et al. 2006).

The EBH run improves the physical model in the simulations in a different direction as it includes prescriptions for the formation, growth, and feedback of supermassive black holes, and assumes “quasar mode” thermal feedback with Bondi–Hoyle–Lyttleton accretion (Bondi 1952; Bondi & Hoyle 1944; Hoyle & Lyttleton 1939; for more details, see Bellovary et al. 2010; Bonoli et al. 2016), while all the rest of the sub-grid modeling and the cooling is identical to that of Eris. Bonoli et al. (2016) found that, except at $z > 3$, when major mergers occur, gas accretion onto the central supermassive black

hole is always well below the Eddington rate. As a result, radiative feedback from gas accretion is negligible. Radiative feedback affects only the very central region (within 1 kpc), resulting in the suppression of the bulge growth relative to Eris. The small bulge is likely the reason why, at $z < 1$, the disk of EBH is more prone to instabilities, which cause the growth of a strong bar at $z < 0.3$ (more details can be found in Spinoso et al. 2016).

The Venus simulation employs different initial conditions. The “zoom-in” was initialized using the MUSIC code (Hahn & Abel 2011), rather than with GRAFIC2 (Bertschinger 2001) as in the other cases, which allows a computationally more efficient topological identification of the Lagrangian subvolume for the refinement. Eris and Venus both form at the intersection of four dark matter filaments, albeit their convergence pattern is different. In general, Venus experiences twice as many major mergers as the other runs, with its last major merger (defined as a merger with mass ratio > 0.1 between the two galaxies) occurring at $z = 0.9$, as opposed to $z = 3.1$ in the other runs. While in Eris a central dominant halo assembles very early, in Venus multiple progenitors of comparable mass evolve separately for a long time, with one single halo and its associated galaxy only appearing after the last major merger at $z < 0.9$. The amount of substructure at $z = 0$ is also more abundant in Venus relative to Eris, both in the stellar and in the dark matter component. In particular, a large satellite orbits around the primary galaxy in Venus even at late times, causing a perturbation on the main disk at pericenter passages, the last of which induces perturbations in the structure of the main disk as late as $z = 0.24$.

3. DISK–BULGE DECOMPOSITIONS

We begin by characterizing the structure of our simulated galaxies, decomposing them into disk and bulge components by two complementary methods. We perform the decompositions of the simulated spiral galaxies at the final redshift z_{end} of each simulation ($z_{\text{end}} = 0$ for Eris, EBH, and Venus; $z_{\text{end}} = 0.3$ for E2k), considering the stellar population in the spherical region within 15 comoving kpc of the minimum of the potential well of the galactic halo. We chose this radius as the limit, upon a visual inspection of the extent of the galaxies at $z = z_{\text{end}}$ in the stellar density maps. We do not identify or subtract particles that might belong to a stellar halo. However, we verified that at z_{end} our results are only mildly sensitive to a range of scale-height thresholds. Moreover, even when we consider all stars within the spherical region, most of the stellar halo is not included in our sample since it would extend to larger radii (see Rashkov et al. 2013). Henceforth we use the term bulge to refer to all stellar particles that do not belong

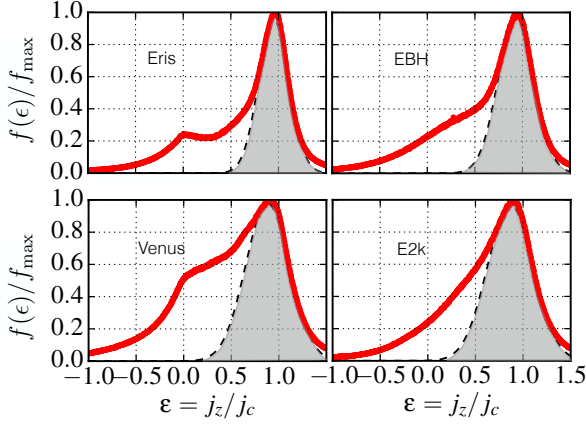


Figure 1. The distribution of the circularity parameter ϵ for stellar particles in a galaxy at the most recent redshift (Eris, EBH, and Venus: $z_{\text{end}} = 0$; E2k: $z_{\text{end}} = 0.3$) is shown with the red line. The black dashed line denotes the Gaussian function obtained as a fit to the distribution right of and about the highest peak in the distribution ($\epsilon \simeq 1$). All particles in the grey-shaded area are assigned to the disk.

to the disk.

The first decomposition method is based on kinematics. In Figure 1, we show four distributions of the circularity parameter $\epsilon \equiv j_z/j_c$, where j_z is the z -component of the angular momentum vector of a stellar particle when the galactic disk lies in the x - y plane, and j_c is the angular momentum of a hypothetical particle at the same location on a circular orbit. One expects the distribution of the circularity parameter, $f(\epsilon) \equiv \Delta N/\Delta\epsilon$ (where ΔN is the number of particles in a circularity bin $\Delta\epsilon$), to show two peaks in a typical spiral galaxy: one at $\epsilon \simeq 1$ for the disk, and one at $\epsilon \simeq 0$ for the bulge.

Using the information about the direction of the total angular momentum of all particles in our sample, we rotate the galaxies to appear face-on in the x - y plane. Then we compute the circularity parameters ϵ of all particles and plot their mass-weighted histogram, which is then normalized to the maximum value of the distribution, f_{max} , as shown in Figure 1 (note that f_{max} represents the most probable value of the circularity).

The total distribution (red line) is the sum of the particles in the disk and bulge components. Some of the galaxies in our sample exhibit non-axisymmetric features such as bars (see the discussions in Bonoli et al. 2016; Guedes et al. 2013); therefore, a standard picture of a Sérsic (1963, 1968) classical bulge and a thin disk (or, equivalently, a sum of two Gaussian distributions of particles peaking at $\epsilon = 0$ and $\epsilon = 1$) is not applicable. However, as seen in Figure 1, all spiral galaxies in our sample have a distinct, close-to-Gaussian distribution for $\epsilon \gtrsim 0.7$. We thus fit a Gaussian function to each of these distributions for all particles with $\epsilon > 0.8$ at

z_{end} to determine the mass and the angular momentum of the disk component, treating the rest as the bulge. The fits are shown with black dashed lines in Figure 1.

To determine the bulge-to-disk (B/D) and bulge-to-total (B/T) ratios, we first denote the Gaussian fit to the disk as $g(\epsilon)$ and define the weighting function $w \equiv [f(\epsilon) - g(\epsilon)]/f(\epsilon)$, which is the fraction of stellar particles in each bin of the circularity histogram assigned to the bulge. The B/D ratios can be determined from

$$D = \int_{\epsilon_{\text{min}}}^{\epsilon_{\text{max}}} [1 - w(\epsilon)] M(\epsilon) d\epsilon, \quad (1)$$

$$B = \int_{\epsilon_{\text{min}}}^{\epsilon_{\text{max}}} w(\epsilon) M(\epsilon) d\epsilon, \quad (2)$$

where $M(\epsilon)d\epsilon$ is the mass of all stellar particles with circularity between ϵ and $\epsilon + d\epsilon$. We set $\epsilon_{\text{min}} = -1.5$ and $\epsilon_{\text{max}} = 1.5$. The final values of ratios B/D and B/T (where $B + D = T$) for the $z = z_{\text{end}}$ galaxies determined with this method are $(B/D)_{\text{kin}} = (0.75, 0.51, 0.96, 0.69)$ and $(B/T)_{\text{kin}} = (0.43, 0.34, 0.49, 0.41)$ for Eris, E2k, Venus, and EBH, respectively.

Although our primary method of decomposition in this paper is kinematic, we also decompose galaxies into disks and bulges photometrically when observed face-on (i.e., parallel to the galactic angular momentum vector). Comparing the results of these two methods may prove useful for determining the uncertainties on the properties derived from the 2D quantities (surface brightness/surface density) vs. those based on the 3D kinematic information.

The mock data, i.e. the surface density profiles for the stars, are calculated as before, for a sphere of $r = 15$ comoving kpc around the centers of galaxies. The fitting function is a combination of the surface density of an exponential disk, Σ_d , and a Sérsic bulge, Σ_b :

$$\Sigma_*(r) = \Sigma_d(r) + \Sigma_b(r) = \Sigma_{d,0} \exp\left(-\frac{r}{R_d}\right) + \Sigma_{b,0} \exp\left[-b_n \left(\frac{r}{R_b}\right)^{1/n} - 1\right], \quad (3)$$

where n is the Sérsic index, R_d and R_b are the scale radii of the two profiles, $\Sigma_{d,0}$ and $\Sigma_{b,0}$ are the central surface densities, and $b_n = 1.9992n - 0.3271$ (Capaccioli 1989). Initially, we determine parameters with a least-squares method, whose best-fit parameters serve as initial guesses to the more sophisticated Metropolis-Hastings Markov chain Monte Carlo method (Metropolis et al. 1953; Hastings 1970) with 10^6 realizations. We have modified this algorithm to accept only those sets of parameters that result in the best fit to our data.

The results are shown in Figure 2, with the best-fit parameters and estimates of bulge-to-disk ratios listed

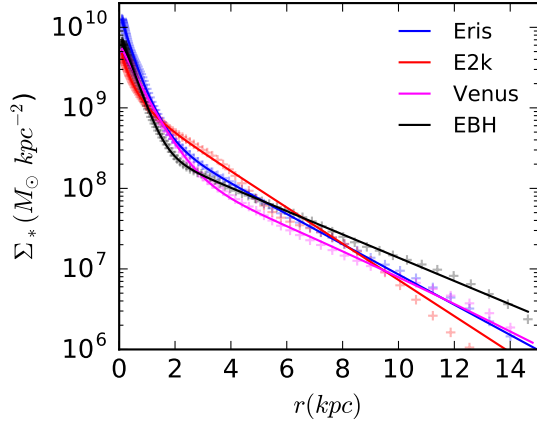


Figure 2. Best-fit functions to the surface-density profiles of our sample of simulated galaxies. The actual surface-density profiles are denoted with crosses, whereas the solid lines show the result of the fitting functions (see Equation 3). The calculations are performed at $z = z_{\text{end}}$.

in Table 2. We computed the bulge-to-disk and bulge-to-total ratios according to $B = \int_0^\infty \Sigma_b(r) 2\pi r dr$ and $D = \int_0^\infty \Sigma_d(r) 2\pi r dr$. When compared to the kinematic estimates, shown again in Table 2, these ratios are in near-perfect agreement for Eris and Venus, but are different for EBH and E2k. Nevertheless, the relative values of $(B/T)_{\text{ph}}$ and $(B/T)_{\text{kin}}$ ratios are consistent between these two methods, i.e. the sequence of runs with increasingly prominent bulge is in both cases: E2k, EBH, Eris, and Venus. Quantitatively, the differences between the two methods are in the 10–30% range for all galaxies except E2k, for which the difference is about a factor of 3.

Each decomposition method has its own limitations, which can affect derivations of the disk-to-bulge ratio and thus measurements of properties associated to it (Abadi et al. 2003; Scannapieco et al. 2010; Guidi et al. 2015). Here, the differences arising in the bulge-to-disk ratios between the photometric and kinematic decomposition methods are strongest for the galaxies with the most prominent bars, which illustrates how the accuracy of the photometric method depends on the type of bulge. Neglecting to separate a bar may lead to an overestimate of the bulge luminosity by 50% (Gadotti 2008). The kinematic method enables us to distinguish between the material of high and low velocity dispersion, while the Sérsic index measures only the curvature of the surface density profile, and does not necessarily describe an object that is truly round or flat in 3D. We note that galaxies with bars might turn out to be morphologically closer to pure disk objects (E2k), *Sc/Sd* galaxies (EBH), or *Sb/Sc* galaxies (Eris), as we do not distinguish bars from bulges in this paper.

The more pronounced difference in the case of E2k,

in which the photometrically defined bulge is much less prominent than the kinematically defined one, highlights a peculiarity of the galaxy with stronger feedback. By low redshift, this galaxy has acquired a stellar component with no clear separation between disk and bulge (see Figure 1). However, when the surface brightness profile is inspected, it is almost a single exponential up to less than 1 kpc from the center, which is typical of very late-type galaxies. The central steepening of the profile inside 1 kpc is highly correlated with the growth of a bar around $z = 1$ and below, as we checked that the inner profile is flatter earlier on. We argue that E2k is essentially a bulgeless¹, barred disk galaxy, or equivalently that the bar makes up for a large fraction of what we identify as the bulge with the photometric method. The tendency of stronger SN feedback to suppress bulge formation (e.g. Hopkins et al. 2014; Keller et al. 2014) is expected since ejective feedback removes low-angular momentum baryons by means of outflows (see Governato et al. 2010; Brooks et al. 2013). The tendency of this galaxy to have a kinematically hotter stellar component is also likely an effect of feedback on the galactic structure, which will be studied in Paper II, but may also reflect the presence of a rather prominent bar which is expected to induce non-circular motions in the stellar component.

We note that, in observations, the quoted B/T ratios are normally obtained by applying the fits and decomposition to the surface brightness profile in a given photometric band, rather than to the surface mass density. Depending on the band, the difference in the relative weight of the bulge and disk can be small or large, with variations of up to a factor of 3 between the B and the I band, depending on the stellar age distributions (Graham & Worley 2008). In general, the bulge, which is composed of an older stellar population than that of the disk, will be fainter in optical bands relative to the disk. Hence our estimates of the relative contribution of the two components, which are based on actual mass density, should be considered as an upper limit. Indeed in the case of Eris, the B/D mass ratio we quote here is higher by a factor of 2 with respect to the I-band B/D found with GALFIT (Peng et al. 2002, 2010) after post-processing with the SUNRISE radiative-transfer code (Jonsson 2006; Jonsson et al. 2010), including dust reddening (Guedes et al. 2011). This supports the notion that, photometrically, E2k is an almost bulgeless galaxy as its B/D would be < 0.1 in optical bands.

¹ Here we define a bulgeless galaxy as one with no detectable extended central component separate from the disk. According to this terminology, a galaxy with a nuclear star cluster such as M33 would be classified as bulgeless.

Table 2. Bulge–disk decomposition parameters

Parameter	Eris	E2k	Venus	EBH
$\Sigma_{d,0}$	6.43×10^8	1.29×10^9	6.12×10^9	3.85×10^8
R_d	2.31	1.93	0.62	2.44
$\Sigma_{b,0}$	2.43×10^9	5.67×10^8	4.97×10^7	1.80×10^9
R_b	0.70	0.70	4.84	0.67
n	1.14	1.35	0.85	0.88
$(B/T)_{\text{ph}}$	0.41	0.11	0.47	0.29
$(B/T)_{\text{kin}}$	0.43	0.34	0.49	0.41
$(B/D)_{\text{ph}}$	0.70	0.13	0.89	0.41
$(B/D)_{\text{kin}}$	0.75	0.51	0.96	0.69
j_d	784.6	762.6	618.4	952.0
j_b	143.0	133.0	101.2	139.3
j_{star}	511.0	547.4	363.3	613.2
j_{gas}	1620.2	916.4	1541.4	1829.6

NOTE—Best-fit values of the parameters of the photometric decomposition (see Equation 3). For a comparison, we also show the results of the kinematic decomposition – $(B/T)_{\text{ph}}$ vs. $(B/T)_{\text{kin}}$ and $(B/D)_{\text{ph}}$ vs. $(B/D)_{\text{kin}}$. We add the values of the specific angular momenta calculated at z_{end} for: j_d – stars in the disk; j_b – stars in the bulge; j_{star} – stars in the disk and bulge; j_{gas} – cold ($T < 10^4$ K) gas in the galaxy. The units of Σ , R , and j are $\text{M}_{\odot} \text{ kpc}^{-2}$, kpc , and $\text{km s}^{-1} \text{ kpc}$, respectively.

We provide a more complete discussion of bulge classification in the Appendix, where we investigate, for each of our simulated galaxies, whether it harbours a pseudobulge or a classical bulge, based on six independent observationally and physically motivated criteria. Some of the methods yield contradictory results, which underlines the importance of examining as many criteria as possible. In essence, we conclude that, at z_{end} , our sample includes composite, classical, pseudo-, and peanut bulges (in Eris, Venus, E2k, and EBH, respectively).

4. J_* – M_* DIAGRAMS

In this section, we use the outcome of the preceding analysis to determine the specific angular momentum of the disk and bulge, as well as that of the overall stellar and gas components. In this way, we can compare the scaling relation between stellar mass and specific angular momentum with those of observed galaxies, as well as study the evolution of such relations from high to low redshifts. This is particularly relevant for the interpretation of the j_* – M_* diagram, which has been proposed as an alternative to the Hubble sequence (Fall 1983; Romanowsky & Fall 2012).

We calculate the specific angular momentum of our sample of galaxies at $z = z_{\text{end}}$ for the entire galaxies and for their separate components. The specific angular momentum vector of particle species k is defined as

$$j_k = \frac{\sum_i m_{k,i} r_{k,i} \times v_{k,i}}{\sum_i m_{k,i}}, \quad (4)$$

where the sums are over each particle i . The particles within each histogram bin (Figure 1) are assigned to a bulge or a disk randomly but in numbers determined by the weight function w defined in Section 3.

The random selection is applied in the interest of simplicity. Although one could additionally classify particles based on their distance from the galactic center or their stellar ages, these simple criteria also bear a degree of arbitrariness – big bulges can extend far from the galactic center (e.g. the Sombrero galaxy), and disks may also include old stars (see Figure 5 of Guedes et al. 2013). This random selection matters most in the transition region of the circularity distribution between the bulge and the disk, thereby it affects a small number of particles. We have verified that assigning the particles in the transition region differently to the bulge and disk (e.g. based on a variety of distance thresholds) does not change the relative weight of disk and bulge by more than 10%.

Our results are set against two samples of observed galaxies. In the left panel of Figure 3, we compare the total specific angular momentum of the stars in our galaxies with a sample of galaxies classified according to their morphologies (Romanowsky & Fall 2012; Fall & Romanowsky 2013). The location of our simulated galaxies on that diagram is well-aligned with the population of observed disk-dominated galaxies. Furthermore, in the right panel, we compare the individual components, i.e. bulges (marked as stars) and disks (marked as squares), with the sample of pure disks and elliptical galaxies. Upon the decomposition, it is evident that all simulated galaxies consist of a disk component with high specific angular momentum and a bulge component with low specific angular momentum. Moreover, in that respect, the disks of the simulated galaxies are in perfect agreement with what is expected of bulgeless galaxies, and their bulges also agree well with what is found for ellipticals. The ratio of specific angular momentum between the corresponding components ranges from 5.5 to 6.8 (see also Table 2).

Morphologically, bulges may be regarded as elliptical galaxies embedded in disks (Kormendy & Kennicutt 2004). In practice, this interpretation might change if a bulge was formed by secular processes (so-called pseudobulge, discussed more in the Appendix). Nevertheless, both the total j_* – M_* diagrams, as well as the dichotomy in the distribution of the specific angular momentum of the components, confirm that the simulated galaxies do not suffer from the overcooling problem or the angular momentum catastrophe (see Section 1), and can be regarded as good laboratories for in-depth studies of the angular momentum evolution.

In what follows, we investigate the evolutionary tracks of the galaxies and their components on the j_* – M_* di-

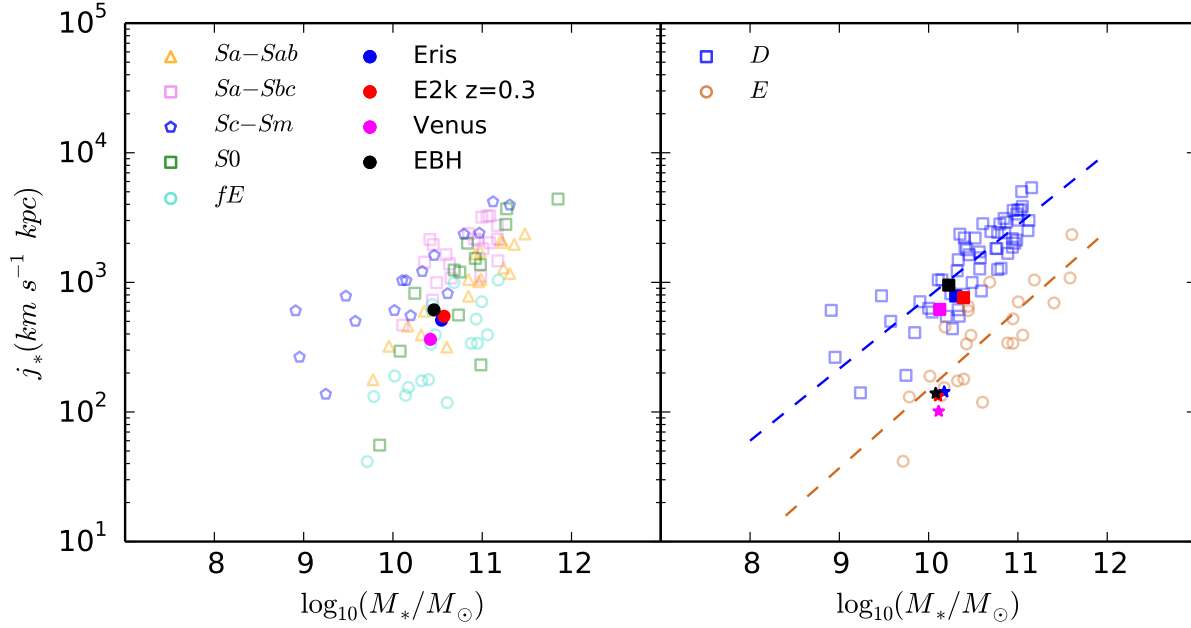


Figure 3. Specific angular momentum–mass (j_* – M_*) diagrams for stars of the simulated galaxies vs. the sample of [Fall & Romanowsky \(2013\)](#). Left. The comparison of the total specific angular momentum of the simulated galaxies with the observed galaxies of various morphological types. Right. Simulated galaxies are kinematically decomposed into disks (filled squares) and bulges (filled stars) and then compared with the subsample of observed pure disk galaxies (D) and ellipticals (E).

agram, which has been proposed as a physically motivated classification scheme alternative to the Hubble sequence, and also address their dependence on the B/T ratio. To do so, first we need to decompose our sample of galaxies at various redshifts. This is a non-trivial task at high redshift, given the complexity of galactic structure, tidal interactions, and frequent mergers. We thus study kinematic diagrams along with the morphology of both gas and stars in order to properly interpret the data. Our results are presented in Figure 4.

In general, upon combining the circularity diagrams with the morphological data-set, in most of the cases the B/T decomposition is straightforward. The same procedure as in the case of the $z = z_{\text{end}}$ galaxies is applied, i.e. based on finding the thin disk in a sphere of 15 comoving kpc encompassing the galaxy. Whenever the fitting of a Gaussian fails (e.g. for both Venus and E2k at $z = 3$), we identify the peak of the circularity distribution that should correspond to the disk, i.e. near $\epsilon = 1$, and then characterize the disk as the ensemble of stars distributed symmetrically around the circularity peak (see for example view *a* for E2k at $z = 3$ in Figure 4).

All galaxies exhibit a similar morphology at $z = 5$, i.e. they appear to be ellipsoids rather than flat extended disks, although in the inner 1–2 kpc a flat disk-like component is already discernible. This and the fact that their circularity distributions peak close to $\epsilon = 0$, lead to the classification of these galaxies as bulge-dominated.

The disky component is approximated by mirroring the distribution around $\epsilon = 1$ as explained above, but with the peculiarity that at this stage there is no peak at high circularity yet.

We caution that the galaxy structure at this redshift might suffer from resolution limitations, as the disk scale length would correspond to only a few gravitational softening lengths at this epoch (disk sizes are expected to be about an order of magnitude smaller simply from the scaling of the halo virial radius with redshift in a Λ -CDM cosmology, see e.g. [Mo et al. 1998](#)). Indeed, recent simulations with much higher resolution (tens of pc) that stop at $z > 5$ do find a prominent rotating disk in halos of masses only a few times larger than ours already at $z = 8$ ([Fiacconi et al. 2016](#); [Pallottini et al. 2016](#)). However, these simulations also find that the disk is thick and turbulent, resulting in $v/\sigma < 2$ for the most part (where v stands for the magnitude of the velocity vector and σ is the total velocity dispersion), which supports the notion that the galaxy would be classified as bulge-dominated based on our criteria. An early phase in which a turbulent gas disk results in a thick primitive stellar disk was already pointed out in [Bird et al. \(2013\)](#).

In Figure 4, the circularity diagrams for Eris and EBH have a dominant rotating disk around $\epsilon \simeq 0.8$, and a secondary peak near $\epsilon = 0$ – 0.1 at $z = 3$. A gaseous disk is evident, and the edge-on view of stars appears flattened, although we witness signs of tidal disturbances from frequent interactions. This suggests that the galac-

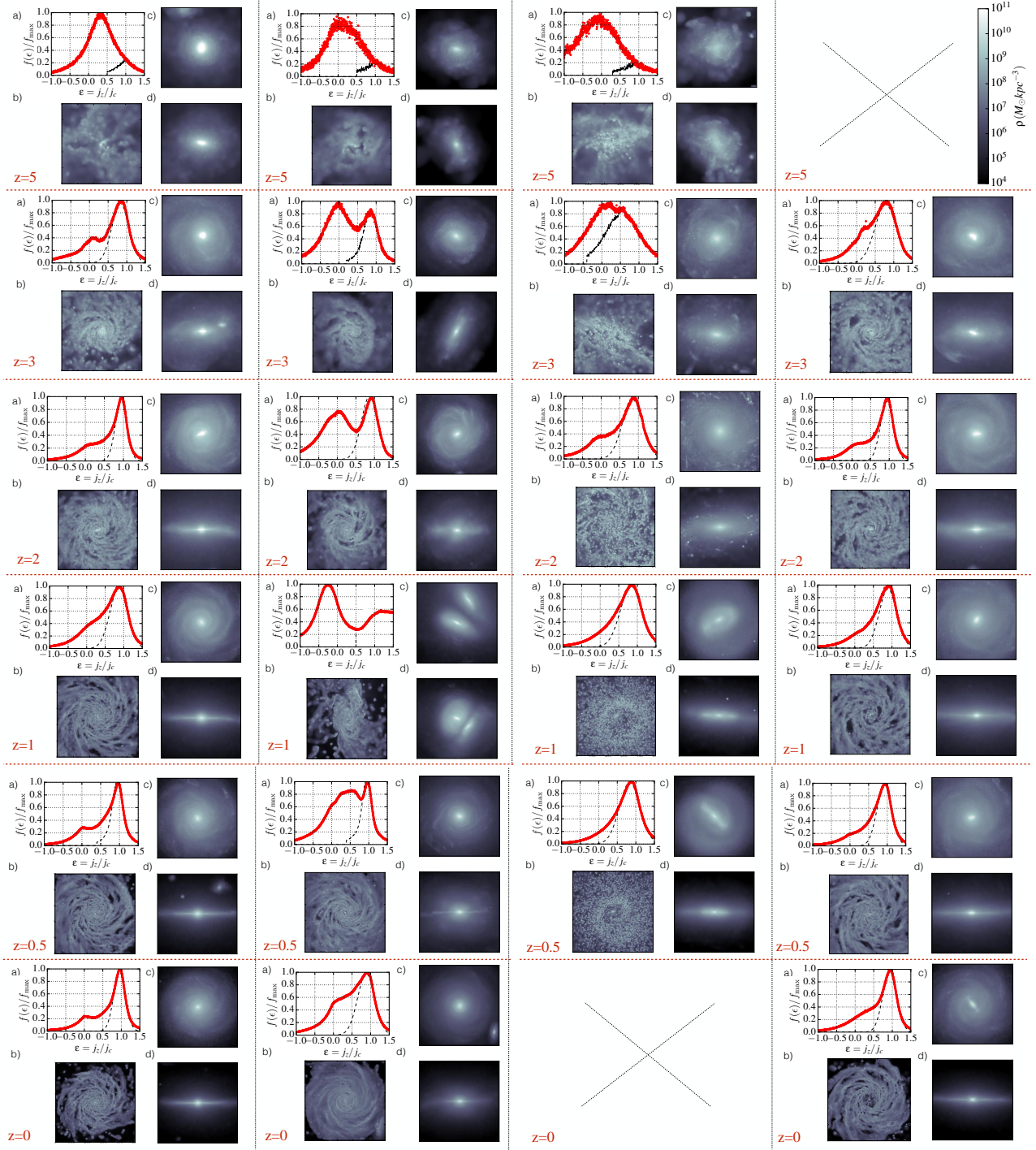


Figure 4. A diagram linking the morphology of a galaxy at various stages of its lifetime with its kinematics. Each column corresponds to a different run – going from left to right: Eris, Venus, E2k, and EBH. Every row shares the same redshift. Crosses are placed whenever an output of a run is missing. Each piece of a matrix contains the following information: a) distribution of the circularity parameter in a galaxy; b) gas density map of a galaxy oriented face-on; c) stellar density map of a galaxy oriented face-on; d) stellar density map of a galaxy oriented edge-on. Total circularity distributions are colored in red, whereas their sub-distributions assigned to the disks are marked in black. Every image has a width of 30 comoving kpc.

tic structure is continually evolving at this epoch and hard to characterize in a simple way. Most of the mass of Venus and E2k has a low circularity parameter peaking at $\epsilon \sim 0$, whereas the secondary peak is lower. The gaseous disk of Venus begins to exhibit a global rotation pattern, in contrast to E2k, which probably reflects the stronger effect of feedback on gas dynamics.

By $z = 2$, all galaxies are already dominated by a thick disk. The face-on stellar density maps reveal spiral structures present in all of them. By this time, Eris, E2k, and EBH have already entered a quiescent phase past the last major merger. Venus, on the other hand, experiences another major merger at $z \sim 1$. In this case, as there are two interacting galaxies at very small separation, the decomposition of the system is somewhat arbitrary. The circularity diagrams show two clear peaks, one at about $\epsilon = 1$ and the other at $\epsilon \simeq -0.2$, thus we choose to cut the distribution at the minimum between the two peaks, i.e. at around $\epsilon = 0.5$. Despite the ongoing stellar merger (views *c-d*), the gaseous disk appears flat (*b*). The spiral structure of E2k vanished, giving way to a prominent bar. The circularity distribution, although strongly asymmetrical, shows only one peak near $\epsilon = 0.8$.

From $z = 0.5$ to 0, the galaxies generally appear similar. The triple-component distribution of Venus settles into one with two components by $z = 0$ with an extended thin disk (in both the gaseous and stellar matter) and a massive stellar bulge. EBH develops a bar (Spinoso et al. 2016) which appears less prominent than the one in E2k in terms of size relative to the disk itself.

With a clear picture of morphological fluctuations in our sample of galaxies, we can proceed to quantify the magnitude of the specific angular momentum vector at various time steps. Figure 5 shows evolutionary tracks in the j - M diagram of our four simulations for their cold gas mass ($T < 10^4$ K, dash-dot line), total stellar mass (solid line), stellar mass in the bulge (dotted line), stellar mass in the disk (dashed black line), and joint total stellar and total cold gas mass (squares) at nine redshifts [$z = (5, 4, 3, 2, 1.5, 1, 0.7, 0.5, 0)$; for EBH and E2k, respectively, the $z = 5$ and $z = 0$ data-points are missing]. The data-points are color-coded with the B/T ratios. The two diagonal dashed lines represent the relationship from Fall & Romanowsky (2013) of $j_* \propto M_*^\alpha$ with $\alpha \sim 2/3$ for disk galaxies (in blue) and ellipticals (in red).

The specific angular momentum of cold gas is substantially higher than that of the stellar component (reaching values of order $10^3 \text{ km s}^{-1} \text{ kpc}$), consistent with observations (Obreschkow & Glazebrook 2014). At all times, disks have a higher specific angular momentum than bulges. At high redshift, all stars have a very low specific angular momentum that evolves below the line

of ellipticals. After $z = 3$, they move over that line and gradually get closer to the line of disks. Eris and EBH evolve on a track with $\alpha = 1.4$, substantially steeper than one with $\alpha = 2/3$. E2k, however, evolves on a track parallel to the one with $\alpha = 2/3$. Venus is a more complex case: it initially evolves approximately along the sequence of elliptical galaxies and, after several fluctuations, its track steepens in the final stages, after the last major merger ($z < 0.7$).

In general, Eris, EBH, and E2k become disk-dominated galaxies past $z = 4$ and evolve into galaxies of different morphological types. Despite the fact that our sample is too small to allow us to generalize, we note that the galaxy with a classical bulge at $z = 0$ has a lower stellar specific angular momentum than those pseudo/composite bulges in Figure 3, and this is also true for the evolutionary tracks of Figure 5. Hence, we argue that the formation path of individual galaxies is also reflected in the final bulge properties.

5. THE GALAXY-HALO CONNECTION

As discussed in the Introduction, theory and observations suggest that the specific angular momenta of galaxies, j_{galaxy} , and those of their dark halos, j_{halo} , are approximately proportional to each other. As the masses of galaxies and halos grow by accretion and merging, their specific angular momenta will also grow (on average). Thus, specific angular momentum is not strictly conserved even in the case $j_{\text{galaxy}} = j_{\text{halo}}$. We expect to learn more about the interplay between the baryons and dark matter in the galaxy formation process by examining the relations between the specific angular momenta of galaxies, that of their disk and bulge components, and that of their dark halos, as well as how these relations evolve with redshift.

We compute the specific angular momenta of stellar disks, stellar bulges, gas at all temperatures, cold gas and dark matter. For each of these components, we define the angular momentum “efficiency” or “retention factor” $\eta \equiv j/j_{\text{halo}}$. Unless otherwise stated, we include all the mass and angular momentum within the virial radius R_{vir} of the halo. It is worth noting that the components of galaxies, as we define them, do not consist of a fixed set of particles; at each redshift, particles are incorporated into or expelled from each component, depending on how galaxies evolve.

Figure 6 shows the evolution of the specific angular momentum of stars (blue dashed line), cold gas ($T < 10^4$ K, black dashed line), total gas (red solid line), stars and cold gas (green solid line), and dark matter (black solid line) in our four simulations. Although the specific angular momentum of each component generally increases with time, as expected, there are noticeable fluctuations at nearly all redshifts. Those tempo-

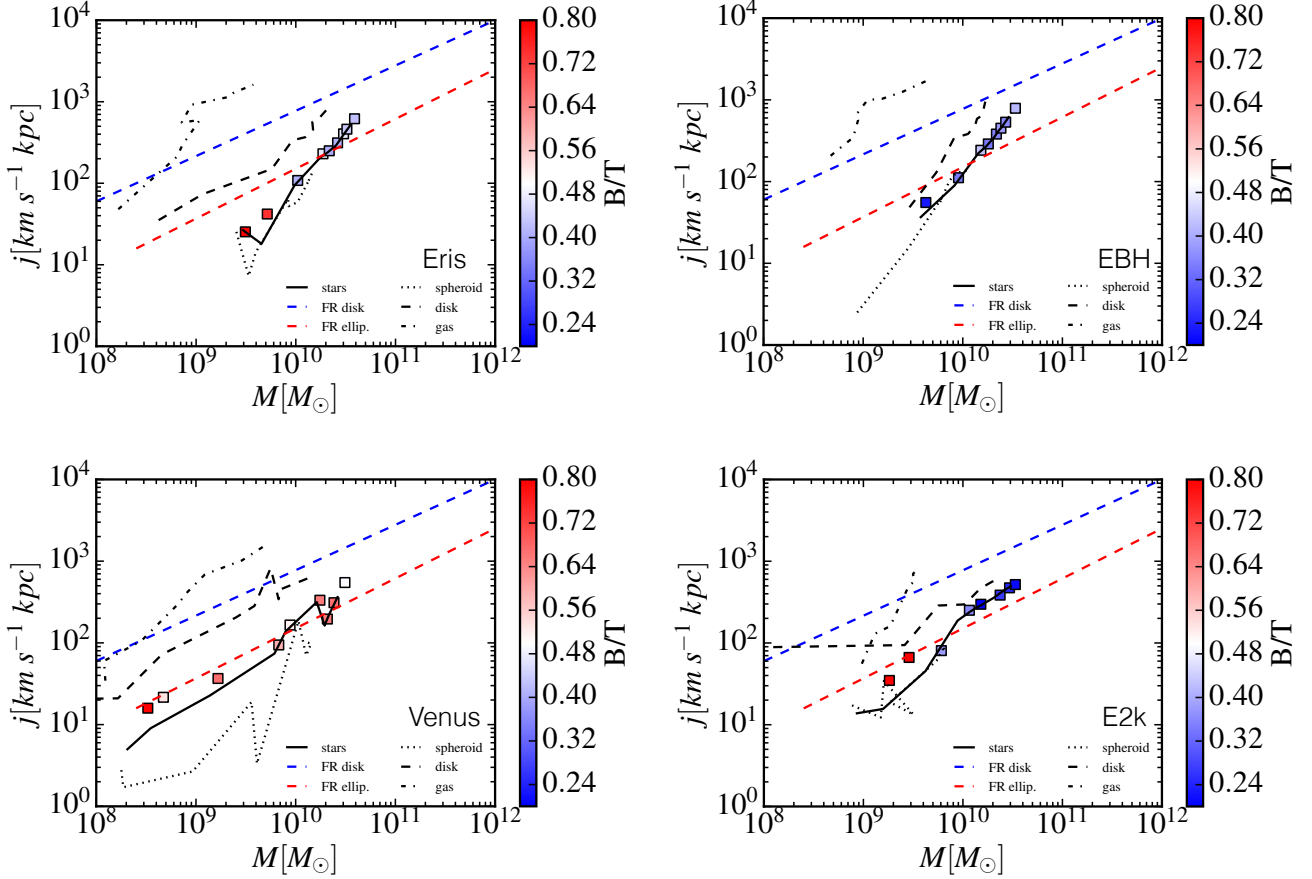


Figure 5. Specific angular momentum evolution of various components on the j – M diagram. The gas considered here is cold (i.e. $T < 10^4$ K). “FR disk” and “FR ellip.” denote best-fit tracks for disk galaxies and ellipticals of [Fall & Romanowsky \(2013\)](#). Each data-point corresponds to the following redshifts (left to right): $z = (5, 4, 3, 2, 1.5, 1, 0.7, 0.5, 0)$ and represents the joint specific angular momentum of cold gas and stars (for EBH and E2k, respectively, the $z = 5$ and $z = 0$ data-points are missing). The color-coding ascribed to the data-points reflects the B/T ratios of the galaxies at a given redshift.

rary gains and losses of the angular momentum per unit mass are stronger before $z = 1$ in all the runs, regardless of the initial conditions. It has been already shown that the angular momentum can be both decrease, e.g. due to torques associated with violent disk instabilities ([Danovich et al. 2015](#)), or increase in galactic fountains if material is ejected for long times and to large radii ([Übler et al. 2014](#)).

[Zavala et al. \(2016\)](#) report a better agreement between the specific angular momentum of the luminous matter and the dark matter within 10% of the virial radius, rather than the full R_{vir} , in their large-volume simulations. We investigate this possibility and show the specific angular momentum for the dark matter within $0.1R_{\text{vir}}$ with a solid grey line in Figure 6. On average, the specific angular momentum of stars evolves closer to that of dark matter within R_{vir} rather than $0.1R_{\text{vir}}$. This holds in all runs, except for Venus in two periods of time: between $z = 2.5$ – 2 and $z = 1$ – 0 . Those two transitions follow sharp changes (reaching an order of mag-

nitude) in the specific angular momentum of all components and are likely associated with major mergers. At those times, the B/T ratios fluctuate (see Figure 5).

As pointed out in Section 4, different initial conditions (active vs. quiet) translate into different locations of the evolutionary tracks of galaxies in Figure 5. As a result, Venus evolves closer to the sequence of elliptical galaxies than the other runs. Since the sequence of ellipticals is offset from that of disks by a factor of ~ 5 , the retention factor for ellipticals is expected to be lower at the same stellar or halo mass ([Fall & Romanowsky 2013](#)). Assuming that the retention factor of pure disks is 0.8–1, one can infer from the j_* – M_* diagram that a retention factor of a galaxy with the same properties as Venus (M_* and j_*) is 0.2. This is close to our simulated value of 0.3.

Figure 7 shows the evolution of the retention factors of the baryonic components of our simulated galaxies. The color-coding is the same as in Figure 6. Evidently, the retention factors exhibit many fluctuations, but, over-

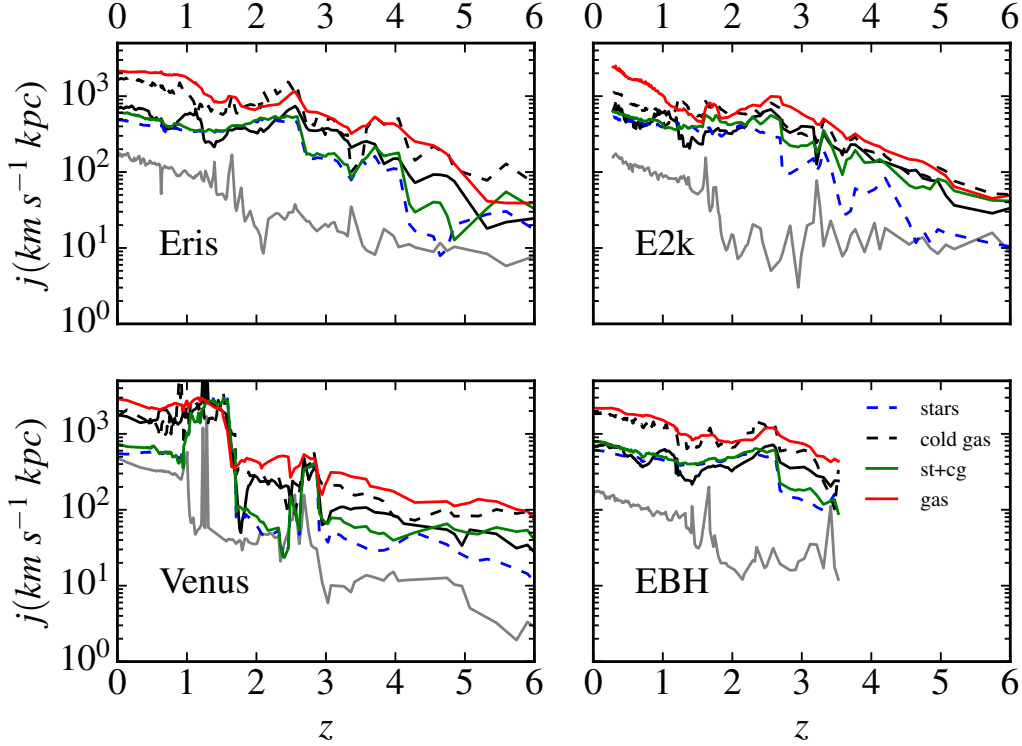


Figure 6. Evolution of the specific angular momentum of the different components of our four galaxies: stars, cold gas ($T < 10^4$ K), all gas, and stars with cold gas within the virial radius as a function of redshift. The solid, black and grey lines represent the specific angular momentum for the dark matter within the virial radius and 10% of the virial radius, respectively. Available data for EBH and E2k exist only for $z < 4$ and $z > 0.3$, respectively.

all, their average evolution is relatively mild. For our sample of simulated galaxies, the stellar retention factor is $\eta \simeq (0.7, 0.8, 0.3, 0.9)$ at z_{end} for Eris, E2k, Venus, and EBH, respectively. The total gas and cold gas are endowed with a specific angular momentum that is 2–6 times higher than that of dark matter (the peak of Venus attains $\eta = 10$ in a major merger). Generally, the retention factor of the stellar component is confined within 0.1–2 in all cases, but most of the time below $z = 3$ it does not exceed 1.5 or fall below 0.5 in the runs with a quiet merger history. The stellar retention factors in all runs are remarkably constant below $z = 1$, which agrees with the theoretical prediction $j \propto M^{2/3}$ for dark halos. Interestingly, despite the relatively violent merger history, the retention factor of stars in Venus at $z = 0$ is almost the same as the initial value (at $z = 6$). Low-resolution, large-volume simulations have converged on the value of the retention factor η for galactic disks that is approximately unity (Genel et al. 2015; Pedrosa & Tissera 2015; Teklu et al. 2015). However, Zavala et al. (2016) found somewhat lower values of η for disks, consistent with a better match to the dark matter within $0.1R_{\text{vir}}$ rather than R_{vir} .

Given the high resolution of our simulations, we can use the results of the last two sections – careful decom-

position and the calculation of the specific angular momentum of the stellar components – to characterize the individual stellar retention factors of disks and bulges at the time steps discussed in Section 4. In Figure 8, we present for the first time the evolution of the retention factors for disks and bulges separately over the redshift range $0 < z < 5$. The retention factors for disks are remarkably constant, ranging from 0.3–1.5, over more than 12 Gyr of evolution. This means that the evolution of the rotationally supported component is driven by its dark matter halo, and this relationship is tightest for systems with quiet merger histories. The retention factors for bulges also evolve relatively slowly but are much lower (by factors ~ 6) than those for disks. Taken together, these results indicate that galaxies and their disk and bulge components maintain nearly constant relationships to their dark halos (after smoothing over short-term fluctuations). That is, galaxies and their halos evolve *quasi-homologously*.

The results presented in this section raise questions about the physical mechanisms that affect galactic angular momentum, as well as the reasons for relatively constant proportionality between j_{disk} , j_{bulge} , and j_{halo} . We plan to address these questions in Paper II. It appears that the strength of SN feedback or the presence

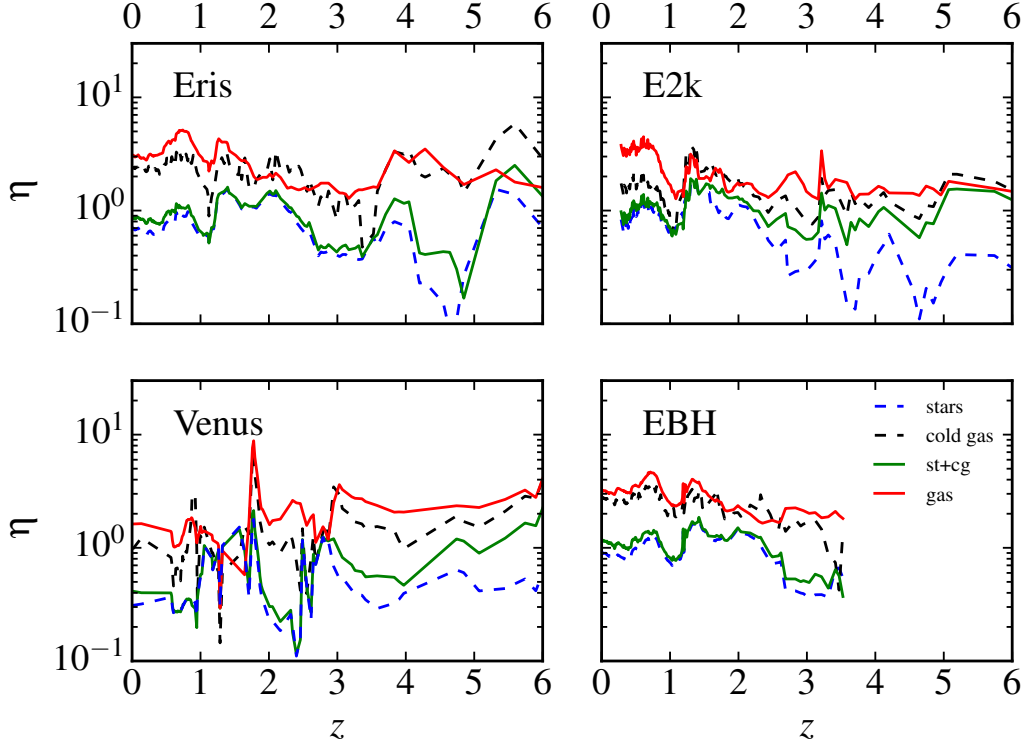


Figure 7. Evolution of the retention factor of the different components of our four galaxies: stars, cold gas ($T < 10^4$ K), all gas, and stars with cold gas within the virial radius at a given redshift. Available data for EBH and E2k exist only for $z < 4$ and $z > 0.3$, respectively.

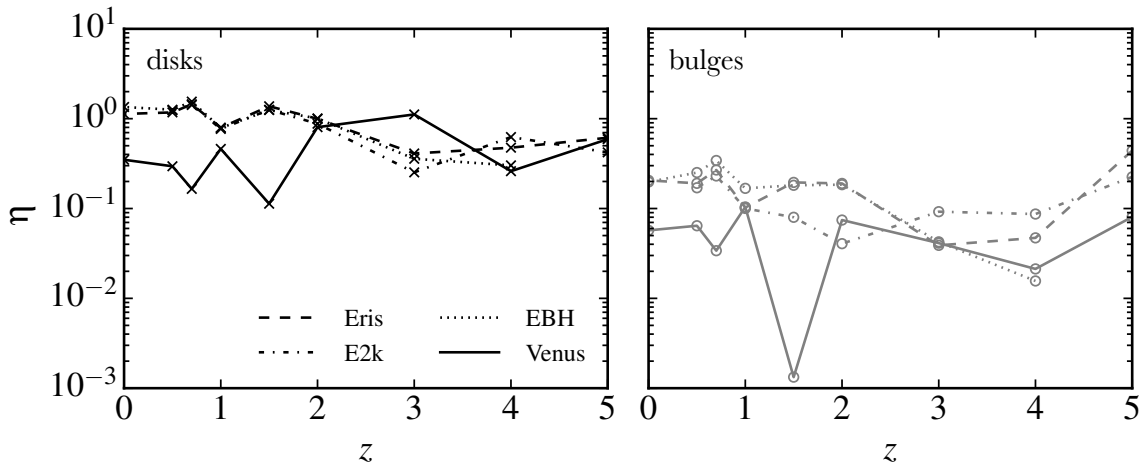


Figure 8. Stellar retention factor for disks (left panel) and bulges (right panel) calculated at a few redshifts, following the kinematic decomposition covered in Section 4. Values for different components are marked with distinctive symbols: disks with crosses and bulges with open circles). The lines denote the interpolations between these data points.

of Milky Way–like AGN (thermal) feedback has less impact on the retention factor within R_{vir} than the merging history driven by the environment in which a galaxy is born (see also the discussion in [Creasey et al. 2015](#), on how galaxy properties depend intimately on their environment). This conclusion, however, comes from a set of simulations employing only one type of a feedback model (blastwave feedback), thus calling for further investigation.

6. CONCLUSIONS

In this paper, we present an analysis of the angular momentum evolution of galaxies residing in Milky Way-sized halos, studying the relation between morphological appearance and kinematics as well as the evolutionary tracks on the j_* – M_* diagrams. We use high-resolution cosmological zoom-in simulations, an approach that is essentially complementary to the large-volume calculations, as we are able to follow the assembly of each object separately and in greater detail. Our sample of simulations comprises runs with varying processes (SN and AGN feedback, as well as different radiative-cooling recipes), different strength of feedback, and different assembly histories for galaxies in halos of identical masses by $z = 0$.

Thanks to the high spatial resolution obtained in our simulations, we can also study the specific angular momentum evolution of the stellar components – the disks and bulges – as the gravitational softening length is an order of magnitude smaller than the characteristic sizes of these components. Additionally, a core part of this paper is devoted to studying the specific angular momentum of gas and dark matter. In what follows, we summarize the main findings of this work and motivate the necessity of a follow-up study.

1. The kinematic decomposition at $z = z_{\text{end}}$ ($z_{\text{end}} = 0$ for Eris, EBH, and Venus; $z_{\text{end}} = 0.3$ for E2k), based on the circularity diagrams as a measure of the kinematics in the plane of the galaxy, and the photometric method yield results that are in perfect agreement for Eris and Venus, slightly deviated in the case of EBH and are off by a factor of 3 for E2k (see Section 3).
2. Our simulated galaxies display a variety of morphological types and lie on the j_* – M_* diagrams with the population of spiral galaxies (Figure 3). When decomposed into disks and bulges, the dichotomy in the specific angular momentum of disks and bulges is reproduced. The disks and bulges of our individual galaxies are separated by a factor 5.5–6.8, in agreement with the findings of [Fall & Romanowsky \(2013\)](#). Our galaxies do not suffer from the angular momentum problem and

are good laboratories for the in-depth studies of the angular momentum evolution.

3. We present time-dependent diagrams (Figure 4) that reveal correlations between the morphological appearance and the stellar kinematics of simulated galaxies, indicating that the latter can be predicted to some extent from the former.
4. We inspected evolutionary tracks of individual galaxies on a physically-motivated equivalent of the Hubble sequence (j_* – M_* diagram). On average, galaxies evolve on straight lines past major mergers on the j_* – M_* diagram (Figure 5). Eris and EBH evolve on tracks with $\alpha = 1.4$ (where $\alpha = d \log j_*/d \log M$). This is likely due to the fact that they undergo a series of morphological changes, which in turn modify their B/T ratios. E2k, which exhibits the least variations in this respect, evolves on a track with $\alpha = 2/3$. We argue that galaxies with relatively stable morphologies and secular processes occurring on long timescales move on the $\log j_*$ – $\log M_*$ diagrams along these parallel tracks. Shorter timescale processes could perturb these tracks: frequent mergers may bring these galaxies closer to the tracks of ellipticals, as in the case of Venus. Although our sample of galaxies is too small to test these scaling relations at a fixed redshift, the time-dependent j_* – M_* sequence of a single galaxy with a constant B/T ratio provides an indirect test of this relationship. Recent observational results of [Burkert et al. \(2016\)](#), [Contini et al. \(2016\)](#), and [Huang et al. \(2016\)](#) for disk-dominated galaxies at higher redshifts ($0.2 < z < 3$) lend support to this conjecture.
5. The specific angular momentum of baryons within R_{vir} tracks that of the halo (Figure 7). The value for the total gas and cold gas is 2–6 times higher than that for the dark matter. The stellar retention factor is nearly constant below $z = 1$ and reaches (0.7, 0.8, 0.3, 0.9) at z_{end} in Eris, E2k, Venus, and EBH, respectively. On average, the retention factors of baryonic components evolve weakly with redshift.
6. In general, the specific angular momentum of stars is more consistent with that of the dark matter within the virial radius R_{vir} rather than within $0.1R_{\text{vir}}$ (Figure 6). When a galaxy is disrupted by many major mergers (e.g. Venus), the overall specific angular momentum of stars is lowered, bringing the latter closer to the angular momentum content of a central subregion smaller than the virial region.

7. Galactic disks have nearly constant retention factors of order unity, which implies their specific angular momenta are strongly correlated with those of the dark halos. Exceptions are the phases in which major mergers occur, which is more relevant for Venus than for the other galaxies having quiet merging histories. The retention factors for bulges are in general within a factor of 10 lower for our simulations and are also relatively constant when short-term fluctuations are smoothed out. These results lead to the notion that galaxies and their halos evolve quasi-homologously.

The good agreement with recent observations for angular momenta of low-redshift galaxies (Fall & Romanowsky 2013), and the confirmation of a close connection of the specific angular momentum of dark matter and baryons, in particular of the disks, set the ground-

work for an in-depth study of physical processes driving the evolution of the angular momentum of baryons in these galaxies. This will be presented in a follow-up paper (Paper II). Paper II will also investigate the role of feedback processes and disk instabilities in the angular momentum transport in these galaxies, as well as how the angular momentum of accreted gas evolves up to the point it joins the disk, shedding more light on the relation between the angular momentum of galaxies and that of their host halos.

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APPENDIX

A. NATURE OF BULGES

As mentioned in Section 3, our sample of galaxies might contain classical bulges (C), pseudobulges (P), “peanut” bulges (box), or so-called composite bulges (COMP, e.g. Kormendy & Barentine 2010). Here, we use six tests commonly used in literature (summarized in Table 3) in order to classify the bulges in our simulations, as well as qualify how sensitive the key results of this paper are to this categorization. These criteria are: visual morphology, presence of a bar, Sérsic index, size–mass relation, star formation rate and vertical distribution (for more details, see e.g. Gadotti & dos Anjos 2001; Kormendy & Kennicutt 2004; Gadotti 2009).

Based on pure visual appearance when seen edge-on (see the last row in Figure 4), the bulges of Venus and Eris could be classified as classical bulges at $z = 0$, whereas those of E2k and EBH appear flatter and more disk-like, hence more similar to pseudobulges.

The fact that EBH and E2k are the only two galaxies that host a strong, large-scale bar at low redshift reinforces this distinction (Eris has a bar at higher redshift which weakens and shortens at low redshift, becoming essentially a nuclear bar a few gravitational softenings long).

In terms of the Sérsic index, all galaxies have a rather low n of order 0.8–1.4 at z_{end} (see Table 2). Given that a relatively low Sérsic index photometrically akin to a disk component ($n < 1.5$) is characteristic of pseudobulges, the

values obtained from the photometric decomposition are hardly indicative of classical bulges.

Another criterion utilizing our results of the photometric decomposition is the mass–size relation. In Figure A1, we compare the location of our bulges on the mass–size diagram with the sample of Gadotti (2009), who found unique relations for ellipticals, classical bulges, and pseudobulges. Our galaxies lie on that diagram in the sequence of decreasing importance of the bar with increasing mass, which places E2k in the area of the graph populated by pseudobulges, Eris by classical bulges, and EBH at the intersection of the two. We note that Figure A1 also places Venus high above the line of ellipticals due to a large radius R_b , which was already argued in Section 3 to be unrealistic and likely represents a failure of the photometric method of decomposition. A brighter stellar envelope of Venus suggests that a more prominent stellar halo or a thick disk component contaminates the decomposition. Indeed, restricting the region to a slice of height 1 kpc above and below the disk plane reduces the bulge scale length to ~ 3 kpc.

The fifth criterion, stellar ages, is addressed in Figure A2, where we show the mass distribution of stellar ages of bulges (left) and disks (right). A massive population of young stars (younger than 4 Gyr) in the bulge of E2k is characteristic of still star-forming pseudobulges. At another extreme, Venus experiences a sharp decline in this distribution near 2 Gyr at the level clearly indicating quenching of star formation, which would be consistent with what is expected of a classical bulge. Eris, EBH, and Venus are good examples of galaxies with a rather old bulge and a young disk.

The final criterion, the vertical kinematics, allows for determining whether disks and bulges are kinematically and structurally alike. Hence, in the top and bottom panels of Figure A3, we show the distribution of the vertical velocity and vertical position of the stellar particles, respectively. The components of Eris and EBH have very distinct vertical kinematics, which is an attribute of galaxies with classical bulges. In contrast, E2k exhibits very little distinction, whereas the case of Venus is rather ambiguous, yet closer to the classical picture. In the bottom panel of Figure A3, the vertical position distributions of the bulge and disk particles of EBH are nearly identical, whereas the disk of E2k is thicker than the bulge, both certainly indicative of a pseudobulge. The bulges of Eris and Venus have clearly broader distributions than the disks, hence a significant part of their mass is off-planar.

Table 3. Types of bulges in the simulations at z_{end}

Criterion	Eris	Venus	E2k	EBH
morphology	C	C	P	box
bar	P	C	P	P
Sérsic	P	P	P	P
size–mass	C	C	P	C/P
star formation	C/P	C	P	C/P
vertical distribution	C	C	P	C/P
summary	COMP	C	P	box

NOTE—The results for each classification scheme are labelled as: P – pseudobulge, C – classical bulge, box – peanut bulge, COMP – composite bulge. The “morphology” criterion is based on the surface density stellar maps.

Although the criteria do not always agree on the classification, overall there is more evidence that E2k and EBH have pseudobulges, whereas Venus has a classical bulge. The bulge of Eris, however, appears to be a composite bulge, i.e. a small, star-forming disk-like bulge inside a classical bulge (see the last row of Table 3). Our conclusion categorizing the bulge of Eris as a composite bulge rather than a pseudobulge complements the previous in-depth studies of the bar evolution in that simulation (Guedes et al. 2013).

Interestingly, if we look at the problem from the formation point of view, we can notice an interesting trend in Figure 5, namely that the most classical-like bulge, that of Venus, evolves most differently from its disk, while the most pseudo-like bulge evolves the most similarly to its disk. Hence, our results for bulges shown in Figure 5 verify that classical bulges (Venus) have lower angular momentum than pseudobulges or composite bulges (the remainder). Also, while in Venus there is a clearly different evolutionary track between the stars as a whole and the bulge, these tracks are almost coincident in the other three cases. This reinforces the notion that formation of bulges is reflected in their final bulge properties, although there is enough diversity and scatter in such properties, and in the results of

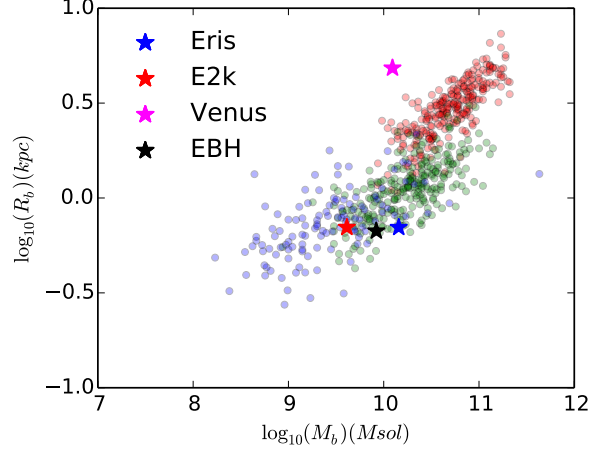


Figure A1. The comparison of bulge scale lengths of our sample of galaxies after the photometric decomposition as a function of the stellar mass of their bulges (stars) vs. the sample of SDSS elliptical galaxies (red circles), classical bulges (green circles), and pseudobulges (blue circles) from [Gadotti \(2009\)](#).

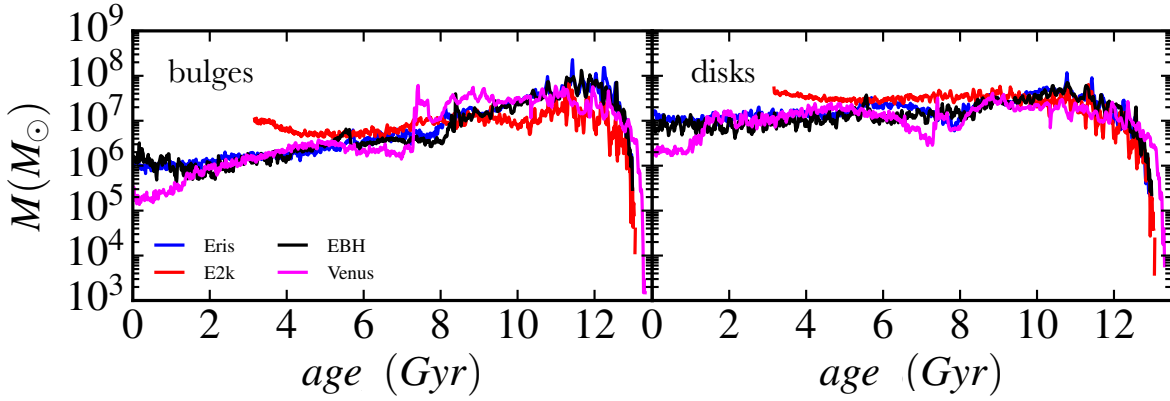


Figure A2. Mass distribution of stellar ages of bulges (left) and disks (right), as would be measured at $z = 0$. Note that the lack of stars of ages lower than 3.5 Gyr in E2k is due to $z_{\text{end}} = 0.3$ of that run.

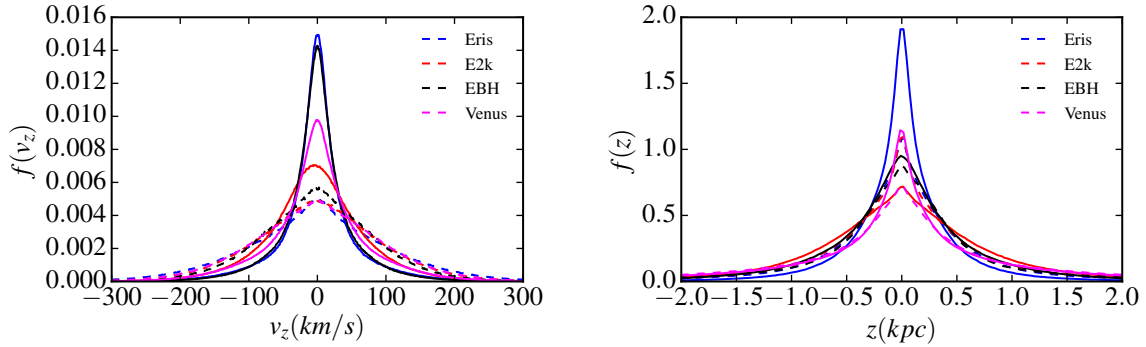


Figure A3. Distribution of the vertical kinematics and vertical locations of the stellar particles in galactic bulges (dashed lines) and disks (solid lines) at $z = z_{\text{end}}$.

different diagnostics, to preclude any rigorous statements.